

How did super-Earths acquire their atmospheres?

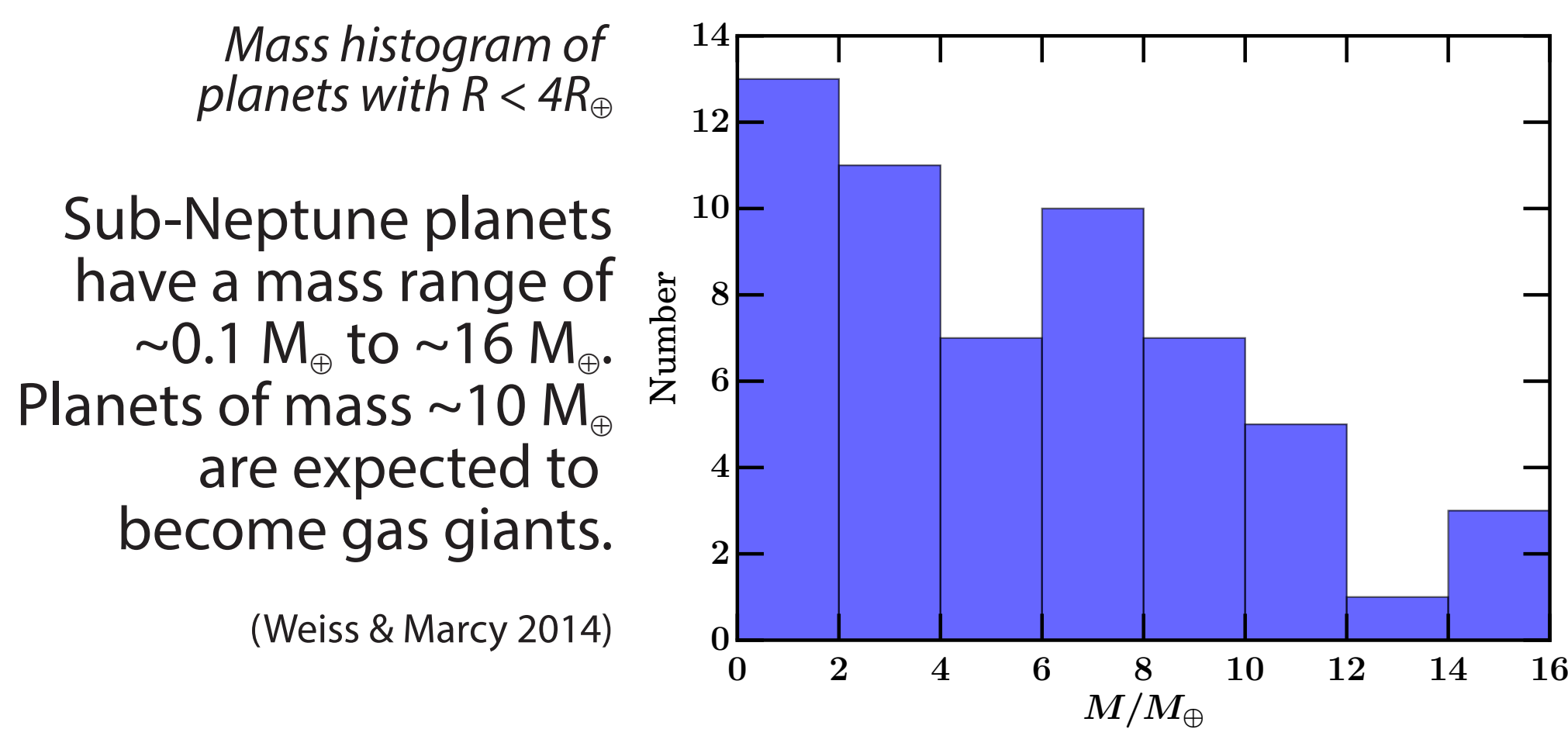
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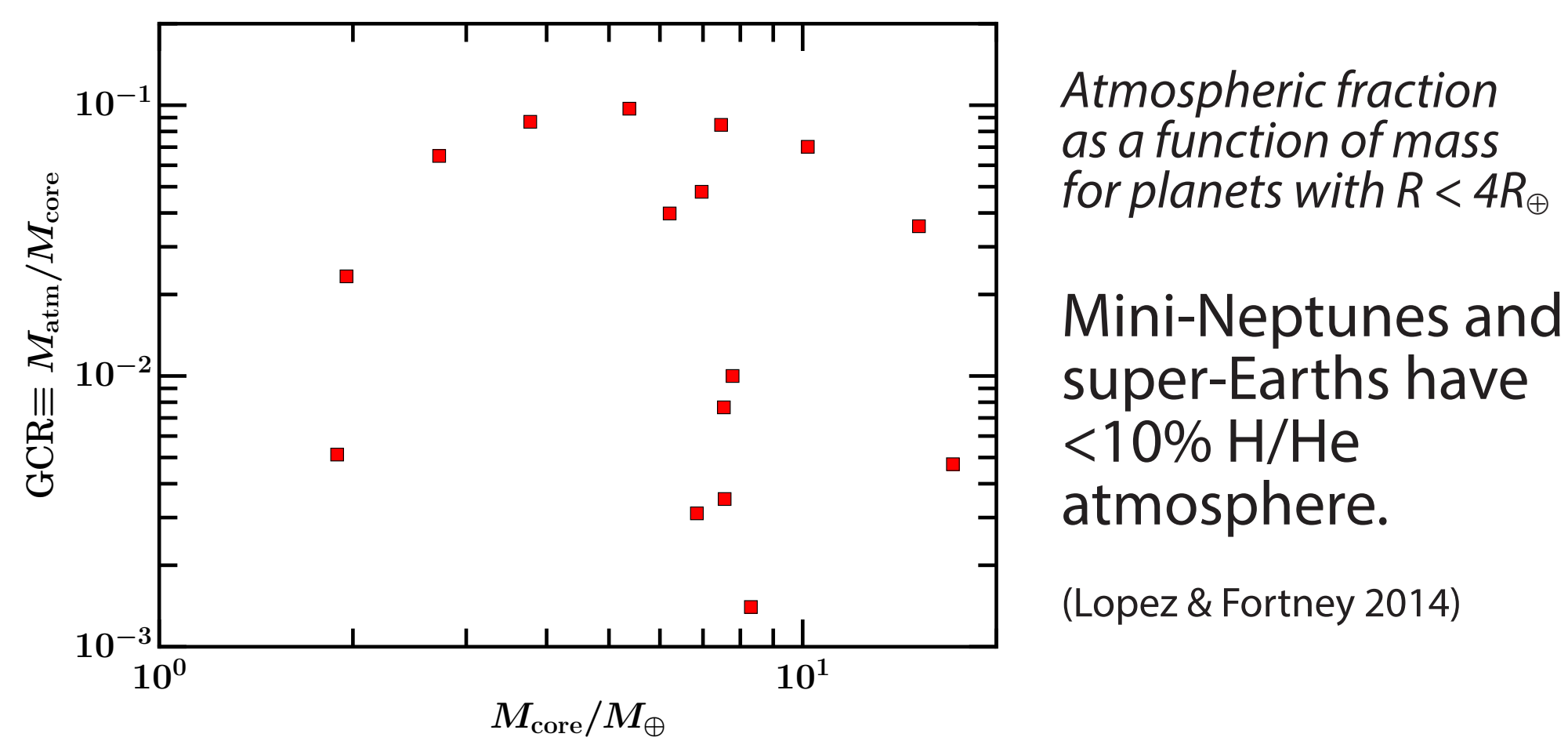
2014 ApJ 797 95



Super-Earths can be as massive as $10 M_{\oplus}$



with a few %wt atmosphere



Why just a few % atmosphere? Can this be explained from primordial gas accretion?

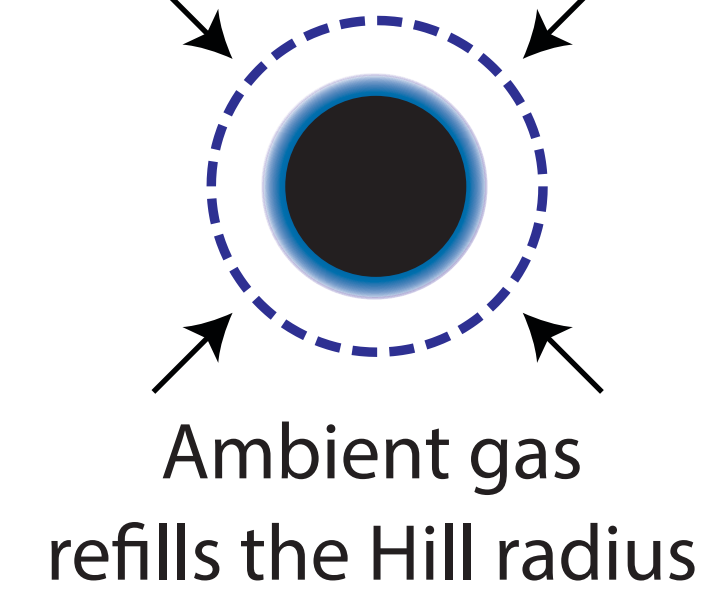
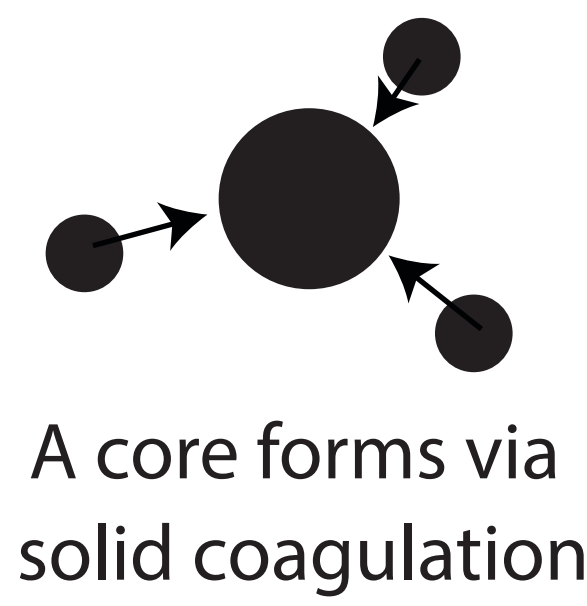
Proto-cores need to have high eccentricity to cross orbits and merge.

Viscous stirring **pumps** eccentricity
Gas dynamical friction **damps** eccentricity

Gas needs to deplete to allow orbit crossing.

Planets accrete gas by cooling their atmosphere

Consider rocks in gaseous nebula



Once $M_{\text{atm}} \sim M_{\text{core}}$, gas self-gravity becomes important. Begins run-away gas accretion.

I ran-away before disk gas depleted!

Disk gas depleted before I had a chance!



¹Smaller of Bondi and Hill radius

Two scenarios of gas acquisition

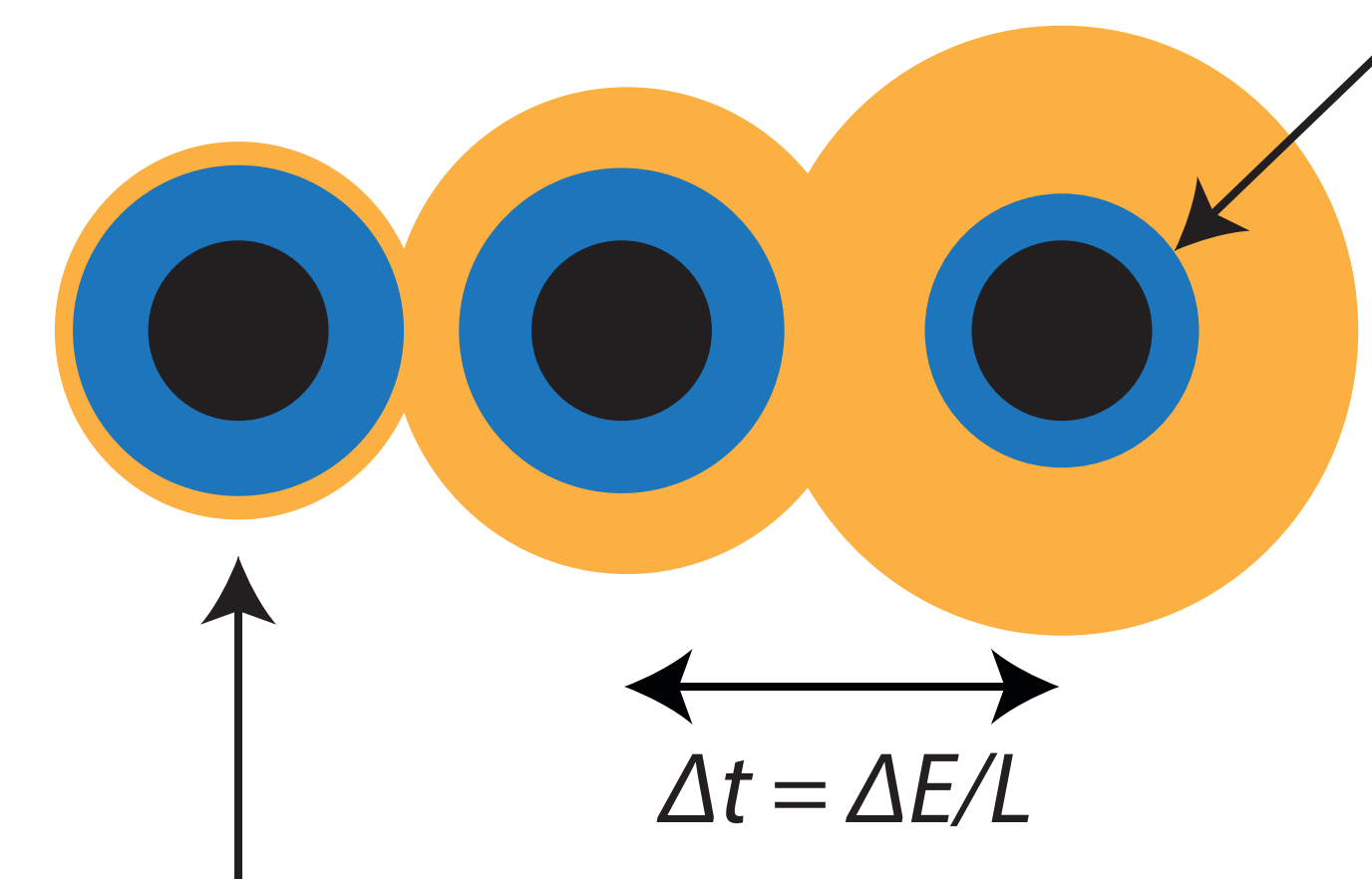
A. Cores form before disk gas depletes

- Cores accrete gas from full nebula
- Avoid gas giant by limiting the rate of cooling

B. Cores form after disk gas depletes

- Cores accrete gas from depleted nebula
- Avoid gas giant by limiting the mass reservoir

Model set-up



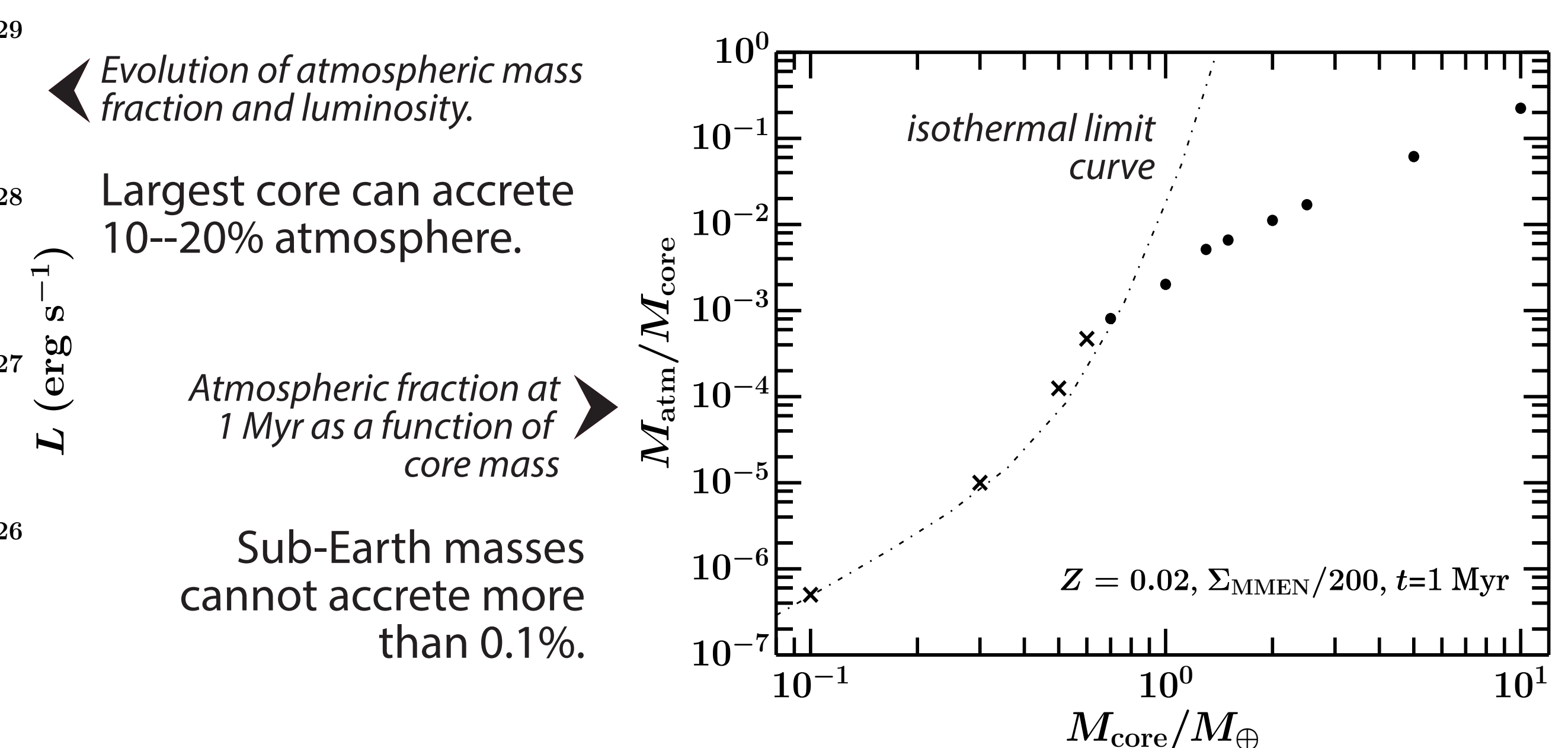
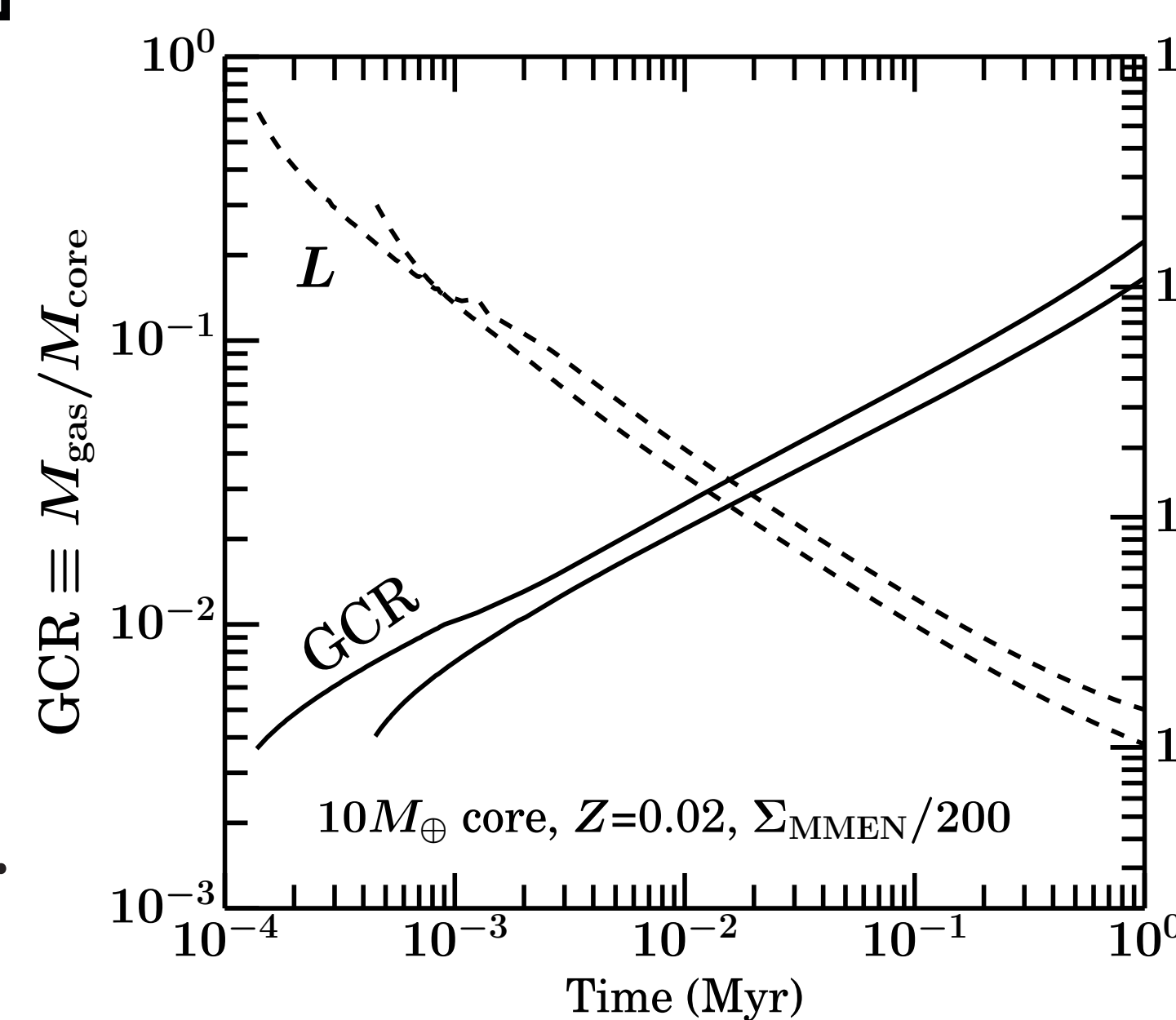
Radiative-convective boundary (RCB); inner convective

- Ideal gas EOS with H_2 , HI , HII , He atoms, metallic atoms
- Opacity with and without grains (Ferguson+2005)

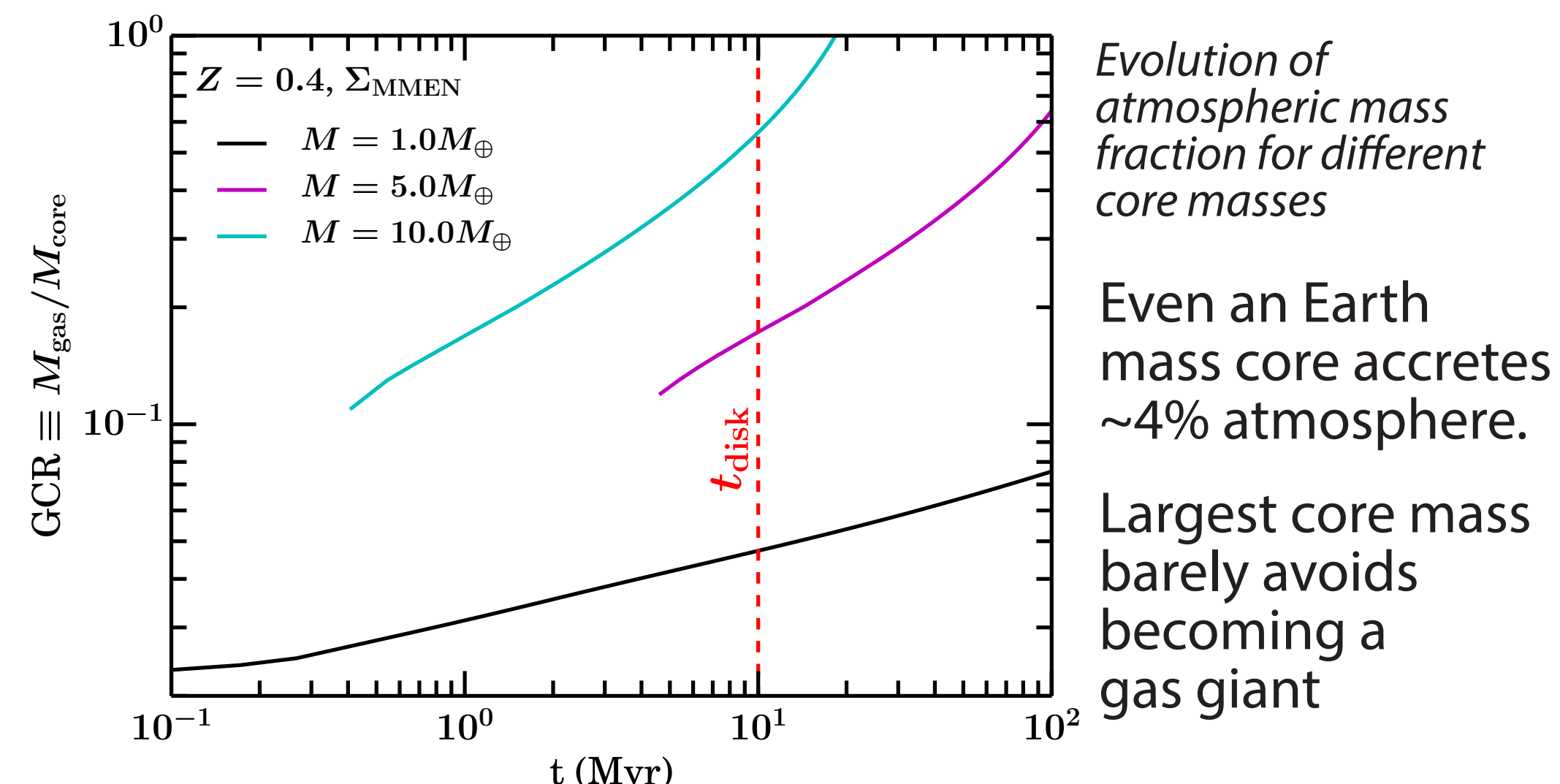
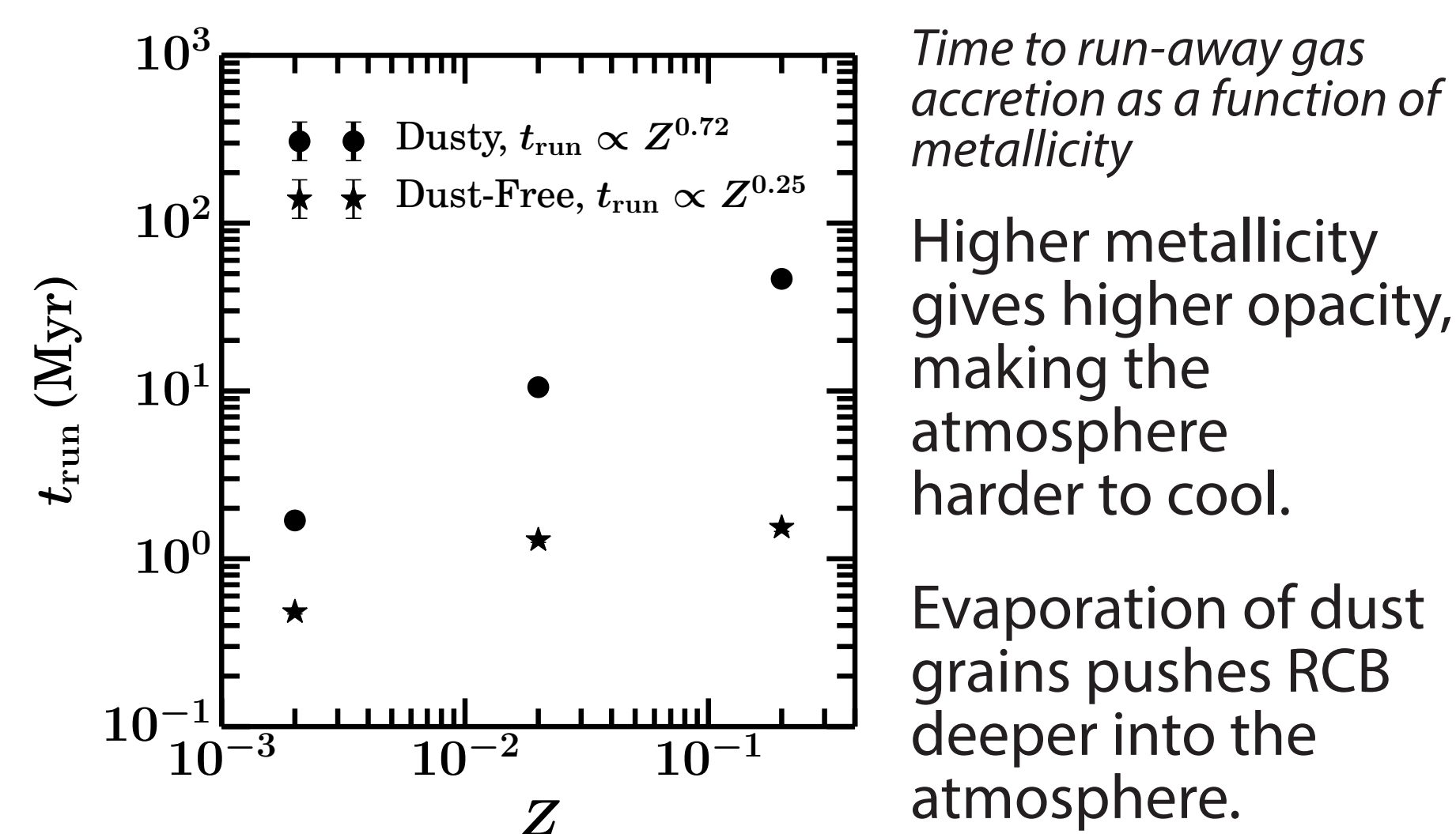
Given a core mass, find L that supports hydrostatic equilibrium

Caveat: no heating

Gas accretion in depleted nebula gives $< 20\%$ atmosphere



Gas accretion in full nebula requires high metallicity



Primordial gas accretion CAN reproduce super-Earth atmospheres!

References

- Ferguson, J. W., Alexander, D. R., et al. 2005, ApJ, 623, 585
- Fressin, F., Torres, G., Charbonneau, D., et al. 2013, ApJ, 766, 81
- Lopez, E. D., & Fortney, J. J. 2014, ApJ, 792, 1
- Weiss, L. M., & Marcy, G. W. 2014, ApJL, 783, L6